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VLBI for Gravity Probe B. V. Proper Motion and Parallax of ...

(SE) in each coordinate Since the proper motion of no bright star was known with such accuracy, we undertook a dedicated program of astrometry to

determine this proper motion for the chosen guide star, IM Pegasi (HR 8703) This star is an RS Canum Venaticorum (RS CVn) spectroscopic binary star, with an orbital period of ~25 days and variable

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Proper Motion of the GP-B Guide Star

- Optically bright (< 55 magnitude)
- Radio detectable (>1 mJy correlated flux density)
- Low declination (near equator to maximize detection of frame dragging and to ease solar power problems)
- No other bright optical star in direct vicinity (to avoid complications in guide-star tracking)

One "Survivor":

VLBI for Gravity Probe B. I. Overview

a bright star could serve this function The orbit selected and subsequently achieved for the mission was also largely dictated by the desire to separate the two relativistic effects from each other and to minimize the vector average over the mission of the nonrelativistic, grav-

arXiv:1107.0970v2 [astro-ph.SR] 8 Jul 2011

Nonetheless, bright disks are detected around $32 \pm 5\%$ of A-type dwarfs (Su et al 2006, Wyatt et al 2003), $16 \pm 28\%$ of solar-type dwarfs (Bryden et al 2006, Trilling et al 2008, Najita The central star is at the origin of the coordinate system The trajectory of the passing star is parabolic, coplanar and